

## Scaling craters in carbonates: Electron paramagnetic resonance analysis of shock damage

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**Abstract.** Carbonate samples from the 8.9-Mt nuclear (near-surface explosion) crater, OAK, and a terrestrial impact crater, Meteor Crater, were analyzed for shock damage using electron paramagnetic resonance (EPR). Samples from below the OAK apparent crater floor were obtained from six boreholes, as well as ejecta recovered from the crater floor. The degree of shock damage in the carbonate material was assessed by comparing the sample spectra to spectra of Solenhofen and Kaibab limestone, which had been shocked to known pressures. Analysis of the OAK Crater borehole samples has identified a thin zone of allocthonous highly shocked (10–13 GPa) carbonate material underneath the apparent crater floor. This ~5- to 15-m-thick zone occurs at a maximum depth of ~125 m below current seafloor at the borehole, sited at the initial position of the OAK explosive, and decreases in depth towards the apparent crater edge. Because this zone of allocthonous shocked rock delineates deformed rock below, and a breccia of mobilized sand and collapse debris above, it appears to outline the transient crater. The transient crater volume inferred in this way is found to be  $3.2 \pm 0.2 \times 10^6 \text{ m}^3$ , which is in good agreement with a volume of  $5.3 \times 10^6 \text{ m}^3$  inferred from gravity scaling of laboratory experiments [Schmidt *et al.*, 1986]. A layer of highly shocked material is also found near the surface outside the crater. The latter material could represent a fallout ejecta layer. The ejecta boulders recovered from the present crater floor experienced a range of shock pressures from ~0 to 15 GPa with the more heavily shocked samples all occurring between radii of 360 and ~600 m. Moreover, the fossil content, lithology, and Sr isotopic composition all demonstrate that the initial position of the bulk of the heavily shocked rock ejecta sampled was originally near surface rock at initial depths in the 32 to 45-m depth (below sea level) range. The EPR technique is also sensitive to prehistoric shock damage. This is demonstrated by our study of shocked Kaibab limestone from the 49,000-year-old Meteor (Barringer) Crater Arizona. We found shock damage present in the  $\beta$  member of the Kaibab Formation exposed in the crater walls corresponding to peak shock stress in the 0.3- to 0.6 GPa range. Carbonate ejecta recovered from within the crater experienced shock pressures of up to 0.6 GPa. Assuming shock damage levels of 0.3 to 0.6 GPa for the lightly shocked carbonate on the walls of the Meteor crater, combined with the shock pressure versus distance model of Moss [1988] and Lamb *et al.* [1991], Meteor Crater impact energies of 2.4 to 8.9 Mt are obtained. This approximately agrees with energies of 3.3 to 7.1 Mt calculated from the crater scaling of Schmidt and Housen [1987].

## Introduction

We analyzed shock deformation in carbonate samples taken from OAK Crater, an 8.9-Mt nuclear near-surface explosion crater located on Enewetak Atoll in the equatorial Pacific Ocean, and Meteor (Barringer) Crater, a terrestrial impact crater located near Flagstaff, Ari-

zona. Electron paramagnetic resonance (EPR) spectrometry was used to determine the peak shock stress experienced by carbonate samples from both craters. Laboratory calibration experiments were used to relate degradation of EPR spectra to shock pressure levels. The OAK results are based on EPR spectra from 136 samples taken from six boreholes within and around the crater and 18 ejecta samples recovered from the crater floor. The sampling of Meteor Crater was less comprehensive and consists of eight samples from the crater wall, 12 samples from the continuous ejecta blanket, and eight miscellaneous samples. The Meteor Crater study presented here, although less complete, demonstrates, for the first time, the feasibility of applying the

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EPR analysis technique to geologically young terrestrial impact craters.

Earlier work relating EPR data to shock deformation was developed by *Vizgirda et al.* [1980] using core material from beneath a 18 kt near-surface explosion, CACTUS Crater. Their work demonstrated a linear relationship between shock pressure and the hyperfine splitting in the EPR spectra that originated from  $Mn^{2+}$  substituting as an impurity for  $Ca^{2+}$  in the calcite component of the carbonate. CACTUS is a bowl-shaped crater, and the results of the EPR analysis were fit to a power curve directly relating sample depth to hyperfine splitting. This paper expands upon the previous calibration technique and extends its application to the study of larger, gravity-dominated, craters.

In this study, shock pressures were determined for the carbonate samples from OAK Crater by numerically comparing the sample spectrum to standard spectra of the well-consolidated Solenhofen limestone (rather than coral media), which had been shocked to known pressures in the laboratory. Similarly, several experimentally shocked Kaibab samples were used as a pressure calibration for the Meteor Crater samples. The present approach employs improvements in data reduction methods, as well as greater EPR signal strength.

We believe the EPR technique described here can be a useful tool in the analysis of other terrestrial impact craters. Of the 150 terrestrial craters suspected to be impact structures [Grolier, 1985], at least 27 were formed within carbonate target rock [Grieve, 1982; Short and Bunch, 1968]. Meteor Crater was formed in sedimentary rocks containing the Kaibab limestone, a sandy dolomitic limestone [Shoemaker and Kieffer, 1974]. The Kaibab Formation makes up a significant portion of the crater wall and is strongly represented in the continuous ejecta blanket.

Although both OAK and Meteor craters have similar final apparent diameters (1200 m versus 1100 m, respectively); OAK is much shallower (60 m versus 197 m, respectively). OAK crater is shallower because of extensive postexplosion modification [cf. Henry and Wardlaw, 1987; Tremba et al., 1990].

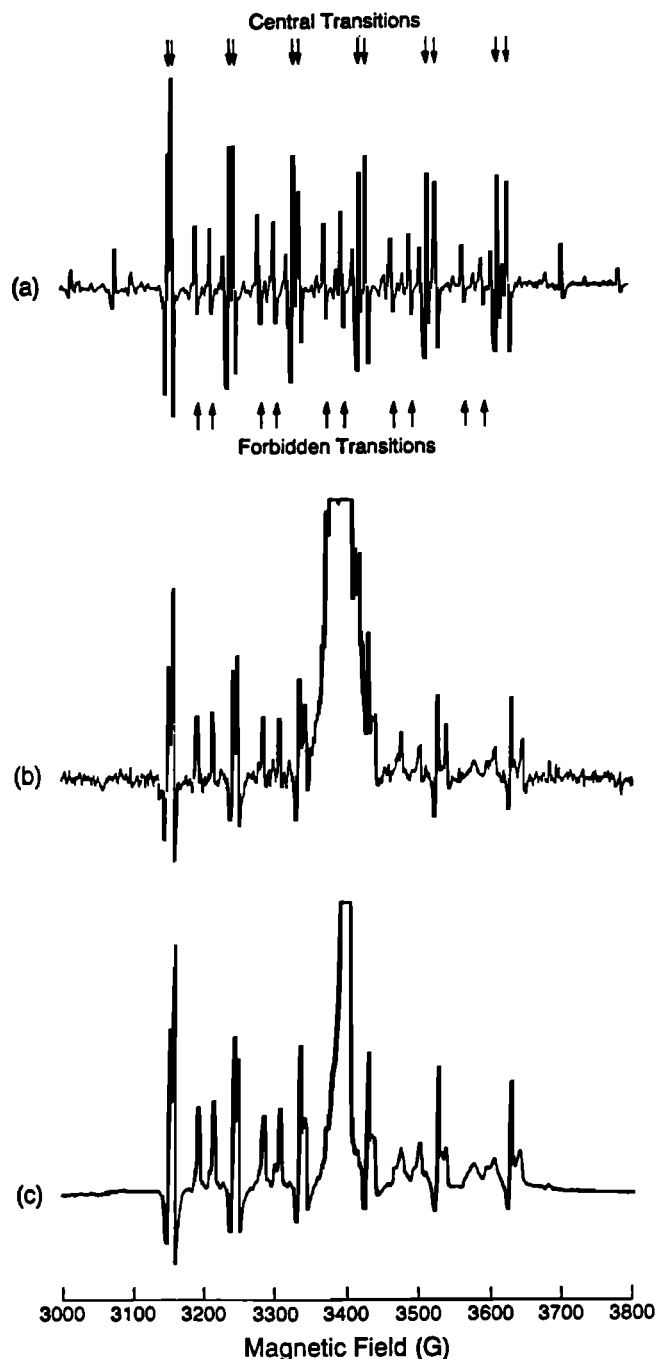
In this paper, we describe properties of the EPR spectrum of  $Mn^{2+}$  in carbonate rocks and the details of how the spectra were obtained and analyzed. Then we discuss the specifics of the OAK and Meteor crater sample analysis. These results are used to define the OAK transient crater, and also compared to volumes obtained with Schmidt et al. scaling laws. In the case of Meteor Crater, we use the peak shock pressure experienced by rocks from the crater walls to infer impact energy and compare this to the values obtained using the scaling law of Schmidt and Housen [1987], as well as the results of previous numerical simulations.

## The EPR Spectrum of $Mn^{2+}$ in Carbonates

### Powdered Calcite

The EPR (also referred to as electron spin resonance, or ESR) spectrum of powdered calcium carbon-

ate,  $CaCO_3$ , is a result of  $Mn^{2+}$  substituting for  $Ca^{2+}$  in a single site in the crystal lattice. The theory of  $Mn^{2+}$  resonance absorption in single-crystal calcite is described by Hurd et al. [1954]. The  $Mn^{2+}$  ion has an effective spin of  $S = \frac{5}{2}$  and a nuclear spin of  $I = \frac{5}{2}$ ,



**Figure 1.** (a) EPR spectrum of powdered single-crystal calcite. The central transitions are due to  $M_S = \pm\frac{1}{2}$ ,  $\Delta m_I = 0$ , where  $M_S$  and  $m_I$  are the electronic and nuclear magnetic quantum numbers, respectively. The forbidden transitions occur when  $\Delta m_I = 1$ . (b) EPR spectrum, unshocked coral limestone, 50 m bsl, OCT-5 borehole (see Figure 5 for hole location). (c) EPR spectrum, 65 m bsl OCT-5b, shocked to  $4.8 \pm 6$  GPa by OAK explosion. The central, 3400-G peak in Figures 1b and 1c is the result of natural radiation damage [see Vizgirda et al., 1980].

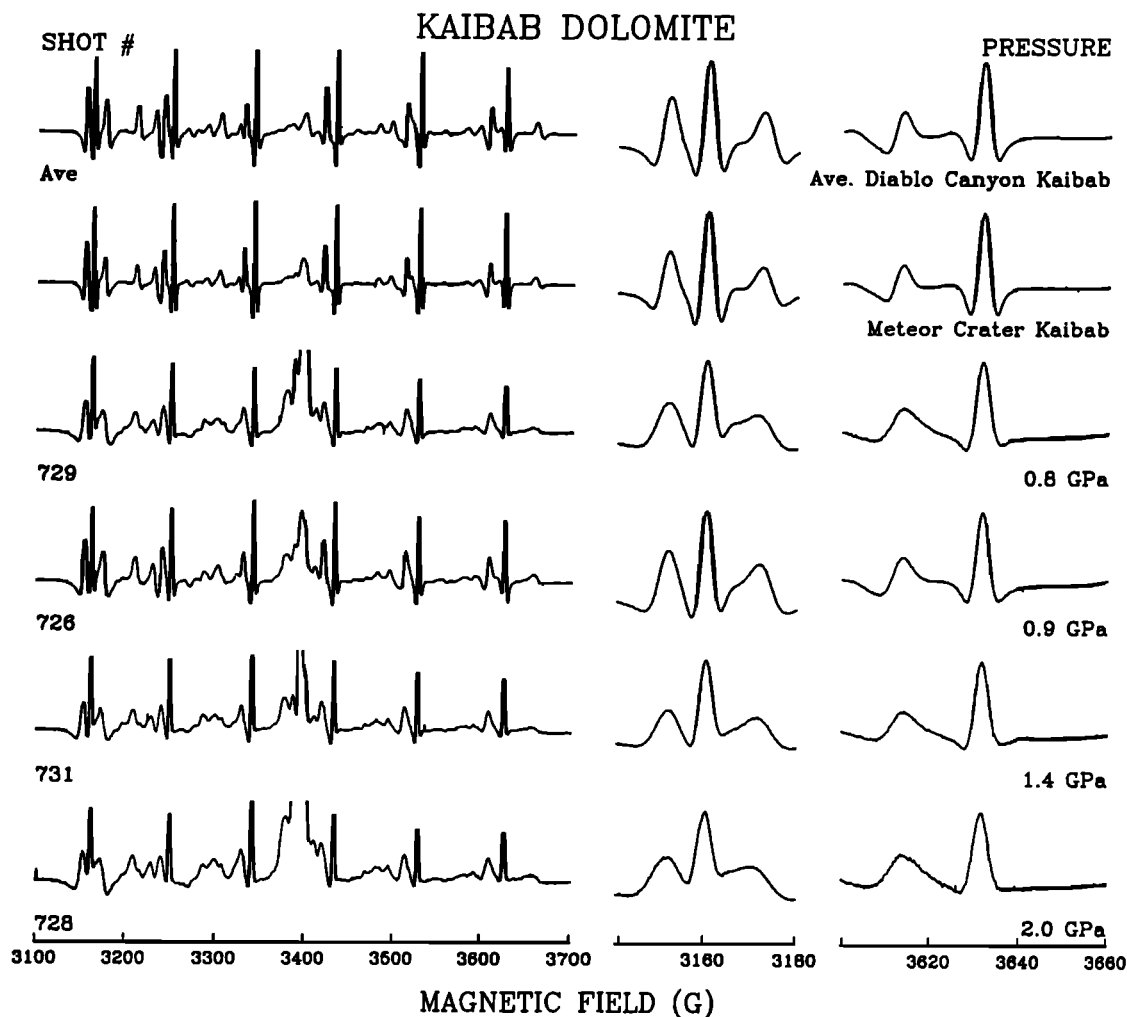
which give rise to a total of 30 transitions. The calcite spectrum exhibits both fine and hyperfine structure but is dominated by the six hyperfine peaks from the central transitions  $M_S = \pm \frac{1}{2}$ , and  $\Delta m_I = 0$ , where  $M_S$  is the electronic magnetic quantum number, and  $m_I$  is the nuclear magnetic quantum number. The hyperfine peak splitting results from the coupling between electronic and nuclear magnetic moments [Hurd *et al.*, 1954]. Another feature of the spectrum is the forbidden transition peaks,  $\Delta m_I = 1$ , which are less prominent and occur in pairs between the central transitions [Mankowitz and Low, 1970]. The spectrum of a powdered sample of single-crystal calcite, Iceland spar, is shown in Figure 1. The central allowed transitions are labeled along with the smaller forbidden transitions. Of particular interest to this study are the two outermost peak doublets at the lowest and highest magnetic field positions of the sextet.

### Coralline Limestone Samples from OAK Crater

The "carbonate" in the OAK crater is a coralline limestone composed of a mixture of calcite and aragonite. Blanchard and Chasteen [1976] have studied both the calcite and the aragonite components of a sea shell using EPR methods. The spectra of  $Mn^{2+}$  in the calcite portion of the shell resemble that of single-crystal calcite; however, no EPR signal was obtained at either room temperature or liquid nitrogen temperatures for powdered samples of the aragonitic portion. Work by Low and Zeira [1972] also describe the lack of an EPR spectrum that is due to  $Mn^{2+}$  in both single crystals and powdered aragonite. They also state that calcite formed from aragonite at low temperatures, such as the case in the Enewetak samples, shows the typical  $Mn^{2+}$  calcite spectrum. Therefore, it is assumed that the spectra obtained in this study (e.g., Figure 2b) are due entirely to the calcite component of the sample.



**Figure 2a.** Comparison of Solenhofen limestone spectra shocked in the laboratory. The first column shows the full spectrum centered at 3400 G, while the second and third columns show the high-resolution spectra of the lowest and highest field components centered at 3160 G and 3630 G, respectively.



**Figure 2b.** Comparison of Kaibab limestone spectra shocked in the laboratory. The first column shows the full spectrum centered at 3400 G, while the second and third columns show the high-resolution spectra of the lowest and highest field components centered at 3160 G and 3630 G, respectively.

### Kaibab Dolomitic Limestone Samples from Meteor Crater

Meteor Crater was formed in sedimentary bedrock containing the Kaibab Formation. The Kaibab consists of fossiliferous marine sandy dolomites, dolomitic limestones, and calcareous sandstones [Shoemaker and Kieffer, 1974]. This formation is exposed on the crater walls and is part of the remaining continuous ejecta blanket. The EPR spectrum of  $\text{Mn}^{2+}$  in dolomite is similar to the OAK carbonates; however, it differs in two important respects. First, the dolomite spectrum is affected by the presence of two different cations,  $\text{Mg}^{2+}$  and  $\text{Ca}^{2+}$ , within the structure [Wildeman, 1970]. Second, because of the difference between the  $\text{Mg}^{2+}$  and  $\text{Mn}^{2+}$  impurities occupy the  $\text{Mg}^{2+}$  sites as well as the  $\text{Ca}^{2+}$  sites. Past measurements indicate a preference for  $\text{Mn}^{2+}$  to occupy the Mg site over the Ca site [Vinokurov et al., 1961; Prissok and Lehmann, 1986].

### Sample Preparation and Spectrometer Measurements

Carbonate samples were ground into a coarse powder and placed into Wilmad 707SQ fused-quartz EPR tubes. EPR spectra were taken at room temperature with a Varian E-Line Century Series spectrometer. The  $\text{Mn}^{2+}$  feature occurs from approximately 3150 to 3650 G, and is centered near 3400 G (where the spectroscopic splitting factor is  $g = 2$ ). The spectrometer was set at 9.56 GHz, microwave power of 20 mW, modulation amplitude of 3.2 G, and time constant of 0.25 s, and operated in the second derivative mode.

### Shock wave Calibration Experiments

The calibration data set is a combination of three series of shock wave experiments. The first series consisted of carbonate samples from CACTUS crater

shocked by us previously [Vizgirda *et al.*, 1980]. The samples were taken from two different depths, 10 and 146 feet (3 and 44.5 m) from the borehole XRU-3 located outside of CACTUS crater on Enewetak Atoll. The principal motivation for reprocessing these samples was to demonstrate that the shock effects observed by Vizgirda *et al.* had remained unchanged. New spectra were taken of each sample, and the results confirmed that the effect of shock on the hyperfine splitting had not altered on the timescale of a decade.

The second series of experiments provided the data used to develop the present pressure calibration technique. The pressure calibration for the coralline Enewetak carbonates was based on six Solenhofen limestone samples shocked to known pressures in the laboratory. Solenhofen limestone was chosen as a calibration standard for the OAK analysis because its EPR spectrum, which was also due to  $Mn^{2+}$  substitution, is orders of magnitude more intense than the spectra from the Enewetak carbonates. The limestone is also more chemically homogeneous, although it is still a polycrystalline material. We note that the carbonate material at Enewetak was saturated, whereas the calibration experiments were conducted on dry samples. What effect water saturation had on the calibrations is uncertain.

The third series of calibration samples consisted of shocked Kaibab dolomitic limestone taken from the walls of Meteor Crater. This calibration was used to determine the degree of shock damage in more heavily shocked samples from Meteor Crater. However, the standard Kaibab comparison spectra were constructed from averages of spectra from five samples taken from Diablo Canyon. Diablo Canyon is  $\approx 10$  crater radii away from the crater, so that the samples should not have been affected by the impact. The Kaibab EPR spectra also have a strong  $Mn^{2+}$  signature of similar amplitude to the Solenhofen limestone.

Cores of sample material, 0.64 cm in diameter, were cut into cylinders 1 cm in length and pressed into stainless steel sample chambers. The rear of the sample chamber cup was enclosed by a stainless steel plug, which was notched to vent any impact-generated gases. The sample chamber was then inserted into a large stainless steel momentum trap and mounted with the

vacuum chamber of the Caltech 40-mm propellant gun apparatus. Lexan projectiles containing flyer plates of aluminum or lexan impacted the target assembly at velocities between 0.8 and 1.6 km/s to yield initial shock pressures of 1.3 to 9.8 GPa and 0.7 to 2.0 GPa for the Solenhofen and Kaibab samples, respectively. Initial shock pressure, rather than final, reverberated shock pressure is quoted because most of the entropy generated by the shock, and hence the shock damage, is associated with the initial shock wave [Lange and Ahrens, 1986].

Shock pressures were calculated using the projectile velocities and the impedance match technique (Table 1) [Stöffler, 1972]. Here  $C_0$  and  $s$  (6) are the usual parameters in a linear shock velocity-particle fit to Hugoniot data [Marsh, 1980]. The average bulk density of the limestone samples was 2.61 g/cm<sup>3</sup>, and the Hugoniot data for Solenhofen limestone were taken from Tyburczy and Ahrens [1986] and Ahrens and Gregson [1964]. A Hugoniot for the alpha member of the Kaibab limestone (2.22 g/cm<sup>3</sup>) was measured by Isbell *et al.* [1966] for shock pressures between 30 and 112 GPa.

### Description of Shocked Solenhofen Limestone Spectra

The spectra of shocked limestone (Figure 2a) have been normalized such that the highest peak of each sample is equal to one. The shocked limestone spectra not only reflect the decrease in the hyperfine splitting observed earlier by Vizgirda *et al.* [1980] in the carbonate spectra, but also reveal that the relative signal strength and width of the two peaks in each hyperfine component vary in a consistent manner with increasing pressure. It is clear from the second and third columns in Figure 2a that the outermost peak in each doublet decreases in relative amplitude and broadens with increasing shock pressure. This alone may be the cause of the decrease in splitting observed with exposure to increasing shock pressure.

The specific behavior of the peaks in the high-field doublet of the carbonate samples from OAK crater is less obvious and becomes difficult to detect in samples subjected to high shock pressures, because of the high aragonite content and poor ordering in the calcite (see

**Table 1.** Shock Wave Data Used in Present Study

Material	Density, g/cm <sup>3</sup>	$C_0$ , km/s	$s$
Lexan*	1.196	2.631	1.295
Aluminum (2024)*	2.784	5.370	1.290
Steel (304)*	7.890	4.580	1.490
Solenhofen limestone†	2.61	3.269	1.796
Kaibab dolomitic limestone‡	2.22	2.10	1.53

\*Data from Marsh [1980]

†Data from Tyburczy and Ahrens [1986]; Ahrens and Gregson [1964]

‡Data from Isbell *et al.* [1966]

Figure 1c). The high-field doublet in heavily shocked samples is not resolvable (as seen in the Solenhofen limestone sample shocked to 9.8 GPa (Figure 2a)). We note that the amplitude of the entire spectrum tends to decrease with increasing shock pressure. This effect is much more obvious in the OAK samples than in the Solenhofen limestone. The inability to resolve the high-field doublet in heavily shocked samples may be in part due to the poor or disordered crystalline nature of the organically produced, and in some cases, diagenetic calcite and Mg calcite (the result of the possible effects of the  $Mg^{2+}$  substitution on the EPR method was not addressed in this study) within the OAK sediments versus the well-crystallized Solenhofen limestone. In addition, the carbonate mineralogy of the shallow OAK samples, i.e., <350-400 ft (106-122 m) below seafloor, is predominantly aragonite (70-80%) (see, *Ristvet and Tremba in Henry and Wardlaw* [1987]); hence the calcite content is low, which also may influence the  $Mn^{2+}$  EPR analyses.

#### Description of Shocked Kaibab Limestone Spectra

Similarly, the series of shocked Kaibab samples (Table 2) is shown in Figure 2b. One obvious difference in the Kaibab spectra is that the low-field component of the sextet is a triplet rather than a doublet. This is not understood in detail. In general, although the sextet of six hyperfine components is still present, each of its components differs from its counterpart in the Solenhofen spectrum. In addition, the forbidden transitions are not as well defined. Despite these differences, the effect of increasing shock pressure appears to have a similar effect on the spectra. Both the low- and the high-field components broaden with increasing shock pressure, and the relative heights of the lesser peaks in each sextet decrease with respect to the highest peak.

#### Pressure Calibration by Differencing Spectra

The previous calibration technique of *Vizgirda et al.* [1980] relied on measuring the separation, in Gauss, of the two peaks of the highest field component of each spectrum. The splitting of the hyperfine peak, HPS, was related to shock pressure,  $P$ , by the relationship

$$HPS(G) = -0.60P(GPa) + 13.85 \quad (\text{high field}).$$

Although the decrease in the splitting of the peaks is most evident in the high-field hyperfine component, the signal strength of this peak is also the lowest. Therefore, as the signal intensity decreases, the error in measuring hyperfine peak splitting increases. The following technique was developed to incorporate the variations in hyperfine splitting as well as relative peak amplitudes and widths. In addition, the analysis will work equally well for the lowest field component of the spectrum, which generally has a higher amplitude than the highest field component.

The standard chosen for the Solenhofen limestone and Enewetak carbonates was single-crystal calcite, while the standard for the Meteor Crater samples was the wall sample used as material for the shock wave experiments. Both high-resolution spectra from each end of the spectrum were used in the comparison. The digital spectra consisted of 1000 amplitude values evenly spaced over a 100-G field range. Both sample and standard spectra were first normalized by the amplitude of the highest peak of the hyperfine spectrum. The sample spectrum was then translated along the magnetic field axis until the position of its highest subpeak coincided with that of the standard spectrum. Next, the absolute value of the difference in amplitude between the two

**Table 2.** Pressure and Integrated Difference (ID) Data for High-Resolution Spectra From Samples Shocked in Laboratory Recovery Experiments

Shot Number	P, GPa	ID, Low Field	ID, High Field
<i>Solenhofen Limestone</i>			
—	0.0	0.0518	0.1187
726	1.3	0.0585	0.1375
727	1.8	0.0709	0.1701
720	2.6	0.0848	0.1924
717	5.3	0.0951	0.2123
718	6.9	0.1086	0.2324
719	9.8	0.1322	0.2654
<i>Kaibab Dolomite</i>			
—	0.0	0.0251	0.0290
729	0.8	0.0999	0.1610
726	0.9	0.0588	0.0454
731	1.4	0.0944	0.1709
728	2.0	0.1334	0.2190

spectra was calculated for each point over the extent of the doublet. Finally, these individual differences were summed to determine a measure of the “likeness” or the “unlikeness” of the sample spectrum to the standard. This number shall be referred to as the integrated difference, or ID, of the sample, which is given analytically by

$$ID = \sum_{i=n_0}^{n_0+N} \frac{|Y_{\text{standard}}(i) - Y_{\text{sample}}(i)|}{N+1}, \quad (1)$$

where  $n_0$  is the index of the amplitude array corresponding to a magnetic field value 20 G below that of the highest peak of the standard spectrum.  $Y_{\text{standard}}(i)$  and  $Y_{\text{sample}}(i)$  are the normalized amplitudes of the standard and sample spectra, respectively, and  $N$  is the number of data points that are integrated. In the case of

the calcite spectra,  $N$  was chosen as 400, corresponding to 40 G. However, the high-field hyperfine peaks in the Kaibab spectra are broader than the limestone spectra, so the ID for these samples was calculated for 60 G, or  $N = 600$ . The error in ID is determined by performing a similar calculation, where  $Y_{\text{sample}}(i)$  are points in the flat baseline signal on either side of the  $\text{Mn}^{2+}$  peak.

Figure 3 illustrates the application of this procedure on two high-field hyperfine spectra from the limestone calibration experiments. The absolute value of the difference between the amplitudes at each point over a 40 G range in magnetic field is shown in Figure 3b. Figures 3c and 3d demonstrate the same technique applied to a limestone sample that has been shocked to 9.8 GPa. The error is determined by using the same scheme to calculate the integrated difference along a

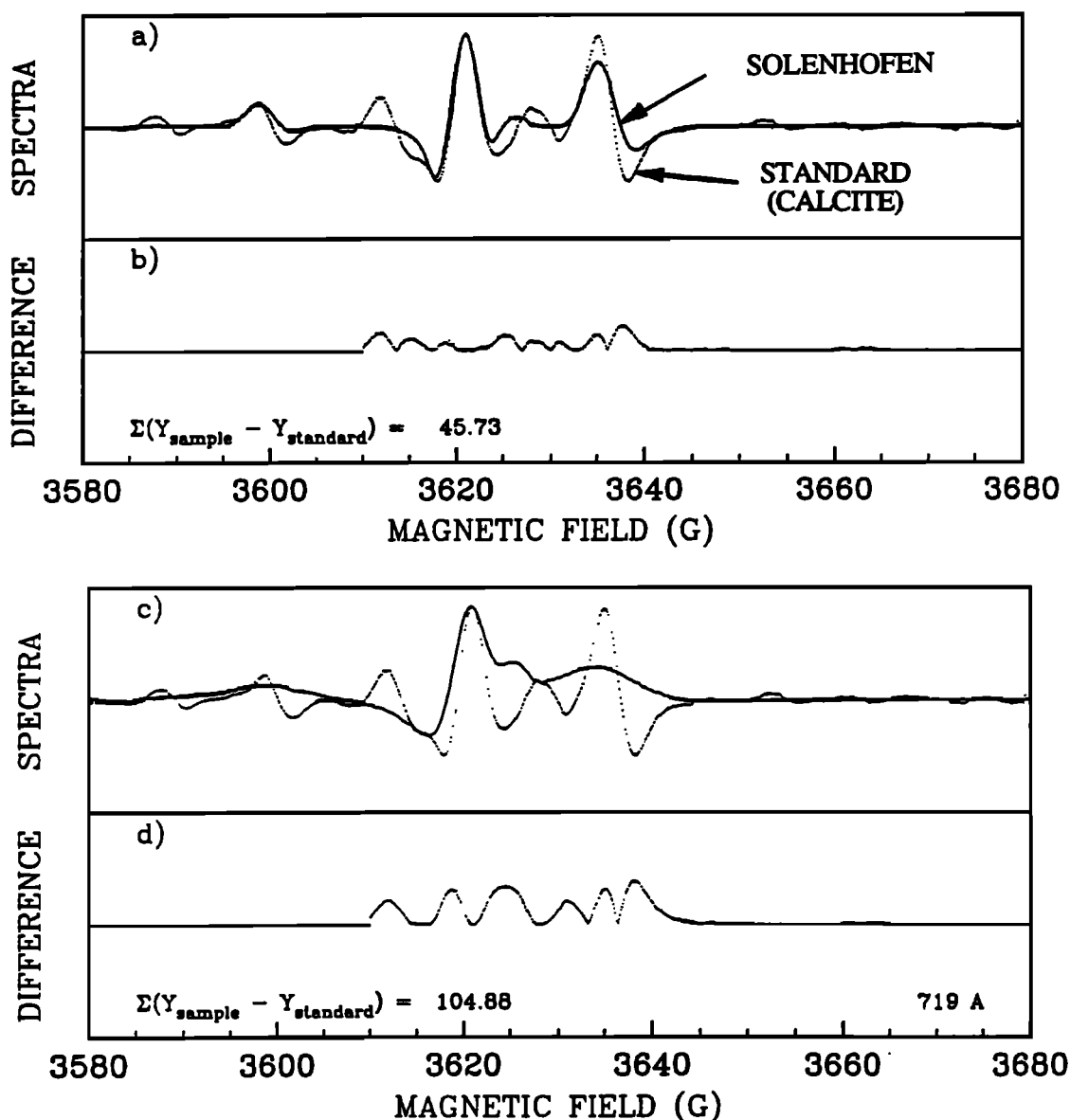


Figure 3. Example of the differencing technique, showing (a) an overlay of the standard (calcite) high-field hyperfine spectra and an unshocked Solenhofen limestone sample, and (b) a plot of the individual absolute differences at each point along the field. Figures 3c and 3d are the same as Figures 3a and 3b for limestone shocked to 9.8 GPa.

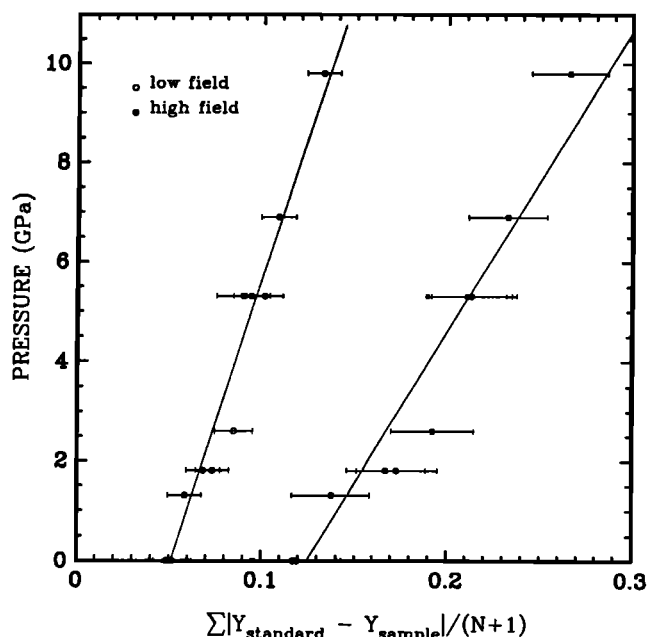
flat portion of the spectrum. This value gives an estimate of the contribution of noise to the ID over the region containing the signal.

The results of these calculations for the limestone calibration experiments are plotted in Figure 4. The ID values are plotted versus pressure for both the low- and the high-field components of the spectrum. To determine the pressure to integrated-difference calibration, a line was fit to each data set, using linear least squares. The resulting equations are

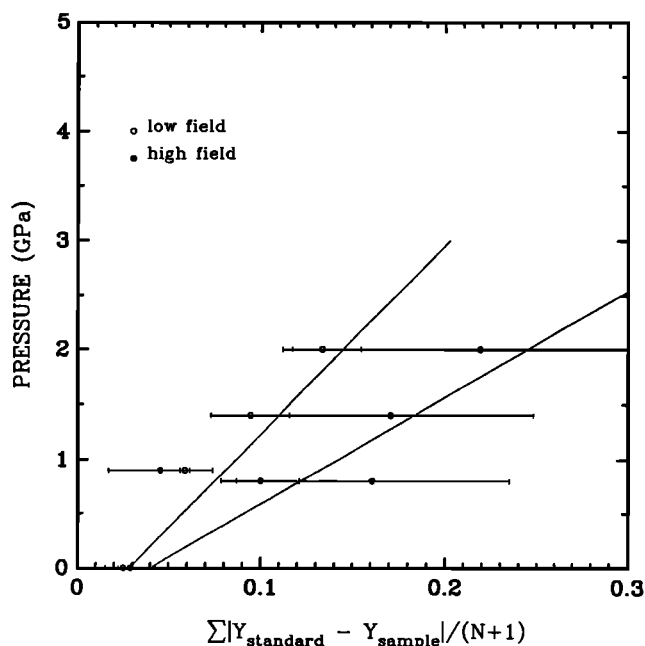
$$P_{\text{Solenhofen}}(\text{GPa}) = 116(\text{ID}) - 5.97 \quad \text{low field;} \quad (2a)$$

$$P_{\text{Solenhofen}}(\text{GPa}) = 61.0(\text{ID}) - 7.60 \quad \text{high field.} \quad (2b)$$

The correlation coefficients for the fit were 0.983 and 0.971, respectively. Table 2 contains a list of the ID results for the limestone experiments. The average ID values are given for shots where several samples were analyzed. Using the calibration curves above, shock pressures were then assigned to the OAK carbonate samples. In general, the carbonate samples have a much weaker EPR signal than the limestone. Therefore, it was necessary to adjust the intercept of the calibration curves to compensate for the average ID value of the unshocked carbonate samples, so that zero pressure is, in fact, obtained. It follows that this method will then assign negative pressures to some samples because the previous adjustment was made to accommodate the "average" background noise. To avoid this obviously unphysical result, and because this technique is



**Figure 4a.** Plot of the summed differences for the low- and high-field components of the Solenhofen limestone samples as a function of shock pressure. The ID value is calculated over a range of 40 G.



**Figure 4b.** Plot of the summed differences for the low- and high-field components of the Kaibab dolomitic limestone samples as a function of shock pressure. The ID value is calculated over a range of 60 G.

not extremely sensitive for low shock damage, all OAK samples with shock pressures calculated to be below 2.0 GPa were classified simply as unshocked. Similarly, the high-pressure cutoff was chosen to be 15 GPa. This is necessary because the rocks were low in calcite content and there are no data for very high shock pressures, and the intensity of the carbonate spectrum is low even at 10 GPa. In most cases, shock pressures were calculated for each sample using both the low- and high-field components of the spectrum. These values were then averaged to determine the final calculated pressure.

Similarly, a calibration curve was determined for the Kaibab experiments (Figure 4b). However, the Kaibab data covered only the low end of the pressure range sampled in the Solenhofen experiments. The fit to the Kaibab data is

$$P_{\text{Kaibab}}(\text{GPa}) = 17.2(\text{ID}) - 0.494 \quad \text{low field,} \quad (3a)$$

$$P_{\text{Kaibab}}(\text{GPa}) = 9.77(\text{ID}) + 0.391 \quad \text{high field,} \quad (3b)$$

where the correlation coefficients for these curves were 0.926 and 0.940, respectively. The slopes of equations (3a) and (3b) are significantly less than the slopes calculated for the Solenhofen limestone experiments. One possible explanation is that this is the result of a sampling bias from computing the calibration for samples in the low-pressure regime. If the Solenhofen calibration was calculated based on only the data up to 2.6 GPa, the slopes would become 71.0 and 31.7 for the low- and high-field components, respectively. Another factor is that the width of each component of the hyperfine sextet of the Kaibab spectrum occurs over a greater



range of magnetic field than its Solenhofen counterpart. This difference could contribute to higher ID values for Kaibab samples relative to Solenhofen samples experiencing similar shock pressures.

## OAK Crater

OAK Crater is located at the northwest portion of the fringing reef at Enewetak Atoll in the Republic of the Marshall Islands. It was formed by an 8.9-Mt explosive device detonated 1.75 m above sea level on a landing craft utility hull approximately 4 m above the lagoon floor ( $1 \text{ Mt} = 4.18 \times 10^{22} \text{ ergs}$ ) [Ristvet *et al.*, 1977]. The present apparent crater has a maximum radius and depth of 875 m and 60 m, respectively. A detailed description of OAK Crater and the geophysical and material studies performed can be found in Folger [1986] and Henry and Wardlaw [1987].

Enewetak Atoll is composed of 42 islands or islets on the Enewetak fringing reef. Some of these islands

are transitory features and represent excess sediment production from the reef that are wind and wave deposited, and in some instances, stabilized by beachrock formation. A summary of their formation is discussed by Ristvet [1987]. The core samples from OAK were obtained through a drilling and coring effort performed from April through July 1985 [Henry *et al.*, 1986]. The rock samples were provided in the form of drill cores 8.6 cm in inside diameter.

A system of geologic zones was defined by Henry and Wardlaw [1987] using the rock nomenclature of Dunham [1962] to describe the stratigraphy beneath the crater. The uppermost Alpha zone consists of late-stage sedimentation and slope failure along with material transported (possibly by hydraulic process) to the surface. The Beta 1 zone involves both early- and late-stage collapse rubble, and also includes some piped material. The Beta 2 zone is referred to as the transition sands. It has limited lateral extent and contains pulverized sand. Beneath the Beta 2 horizon, the rock of the Beta 3

## OAK CRATER BOREHOLES

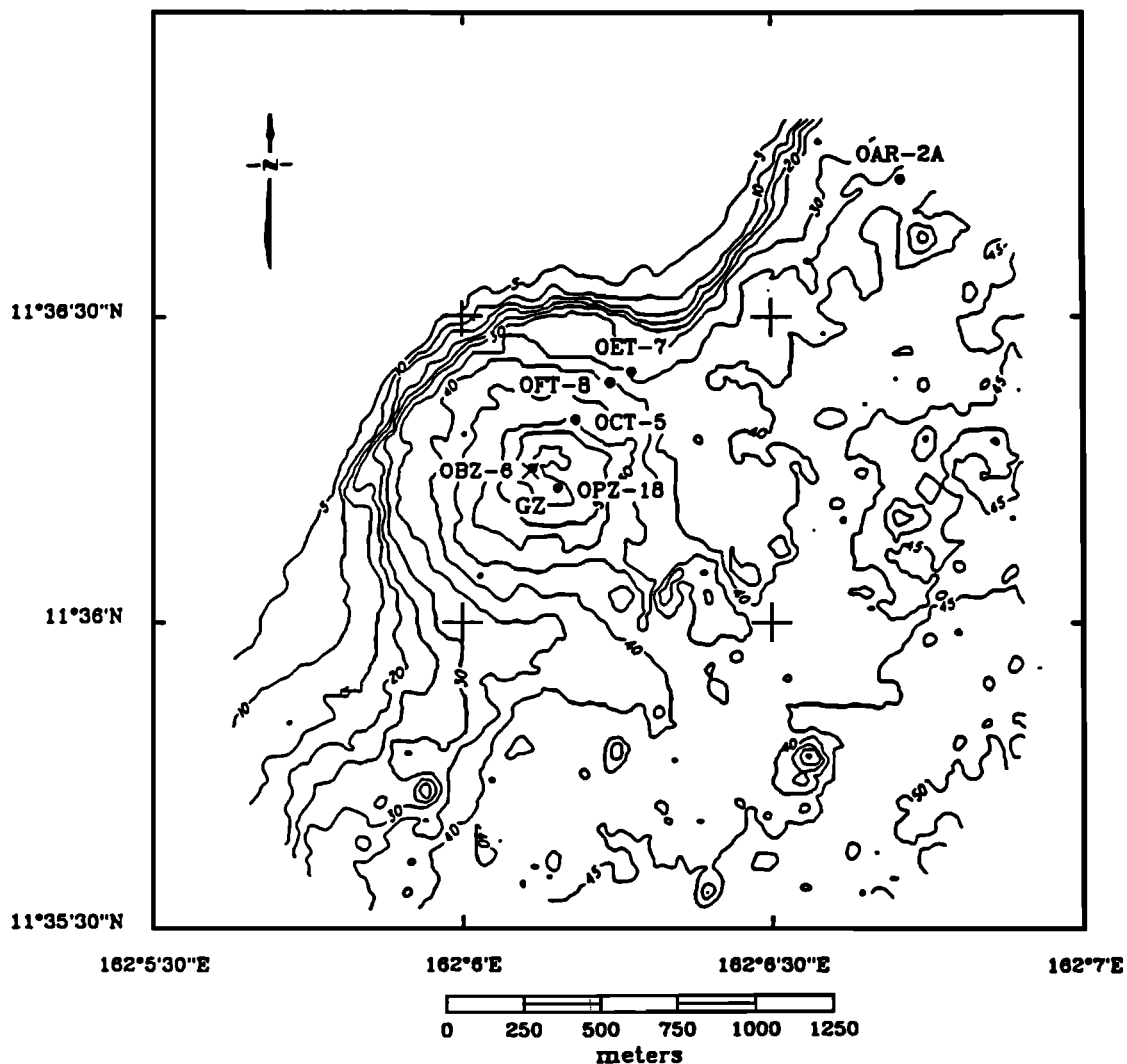


Figure 5. Map of OAK crater [after Folger, 1986] showing the location of the boreholes sampled in this study. Contours are in meters below sea level. OPZ-18 and OPZ-4 are nearly colocated.

and Gamma zones becomes increasingly less fractured and displaced. This system will be used later when discussing the results of the borehole sample analysis.

### Core Sample Selection

The core samples consisted of sediments from the six boreholes OAR-2A, OBZ-4, OCT-5, OET-7, OFT-8, and OPZ-18. These boreholes are located on a north-east crater radial, which is parallel to the reef structure or the geologic strike of the atoll (Figure 5). Core samples are designated by a borehole name and a depth in meters below sea level (m bsl).

The carbonate material from Enewetak is extremely inhomogeneous material containing both calcite and aragonite. As mentioned previously, aragonite does not have a detectable EPR spectrum [Low and Zeira, 1972]; therefore, when possible, samples were selected for high calcite content. For example, those samples containing coral replaced by solution-deposited calcite crystals were preferred because they had stronger EPR signals. We believe these calcite crystals formed during a glacial period when the lagoon was a floor above sea level. The choice of good sample material is important to the extent that it provides a consistent base for analysis, and guards against mistaking a sample with an inherently

poor EPR spectrum as one being heavily shocked. The difference between the two cases can usually be recognized by visual inspection but is more difficult to assess with numerical techniques. In each core, the majority of samples were taken from depths above the Gamma geologic zone, defined by Wardlaw and Henry, [1986]. A more detailed description of each sample and its spectra can be found in Polanskey and Ahrens [1987]. Complete descriptions of the stratigraphy of each borehole are given by Henry *et al.* [1986].

### Results of Core Sample Analysis

The two boreholes located directly below the position of the explosive device, ground zero boreholes OBZ-4 and OPZ-18, were the most heavily sampled cores. A very highly shocked layer of uncemented material was found in OPZ-18 between 121.9 and 126.8 m bsl. This layer was visually distinguishable by the unusual greenish color of the carbonate sand. The shocked zone was broken at 125.7 m bsl by a thin zone of lighter colored material. The location and nature of this shocked material coincide with a zone of Holocene sediments described by Wardlaw and Henry [1986] as a possible example of material that has been injected. The present results are consistent with such a hypothesis, since this

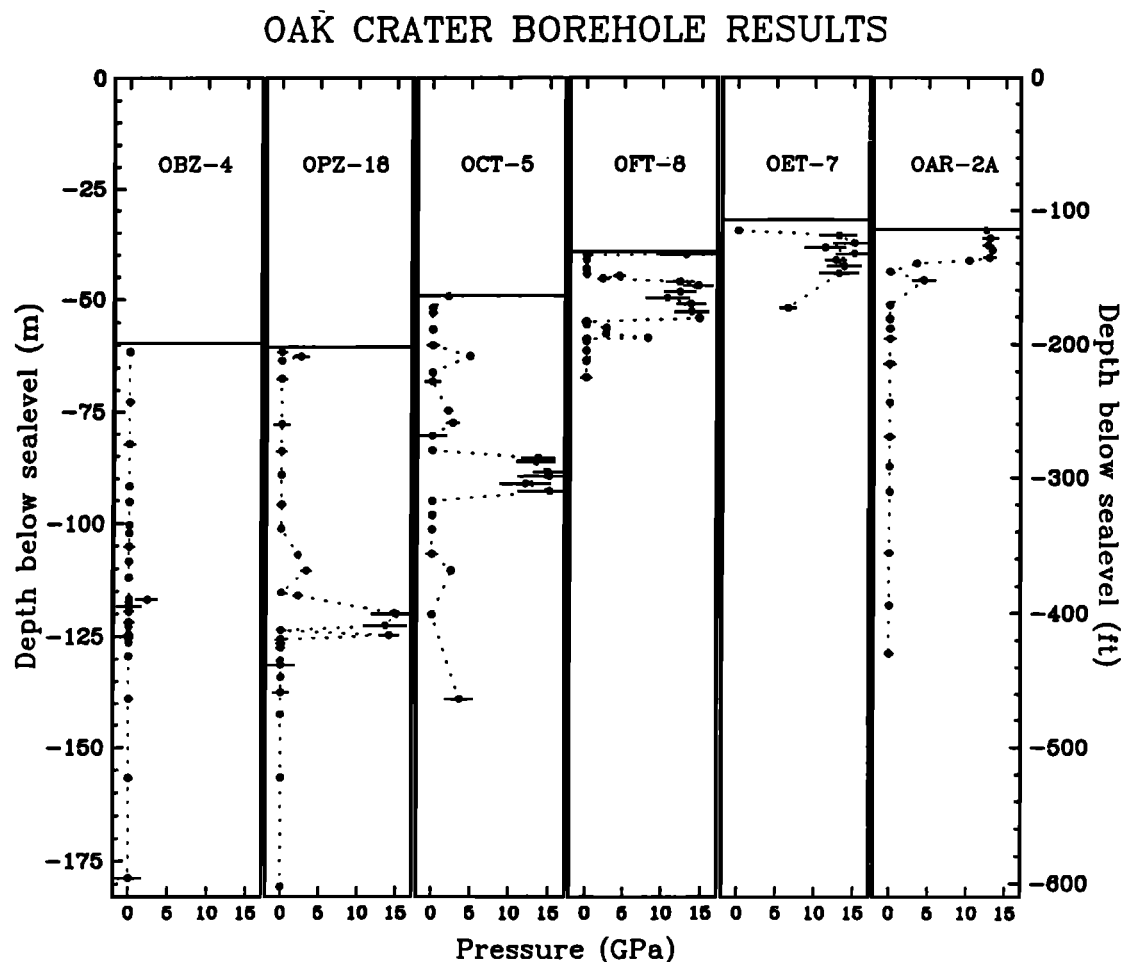
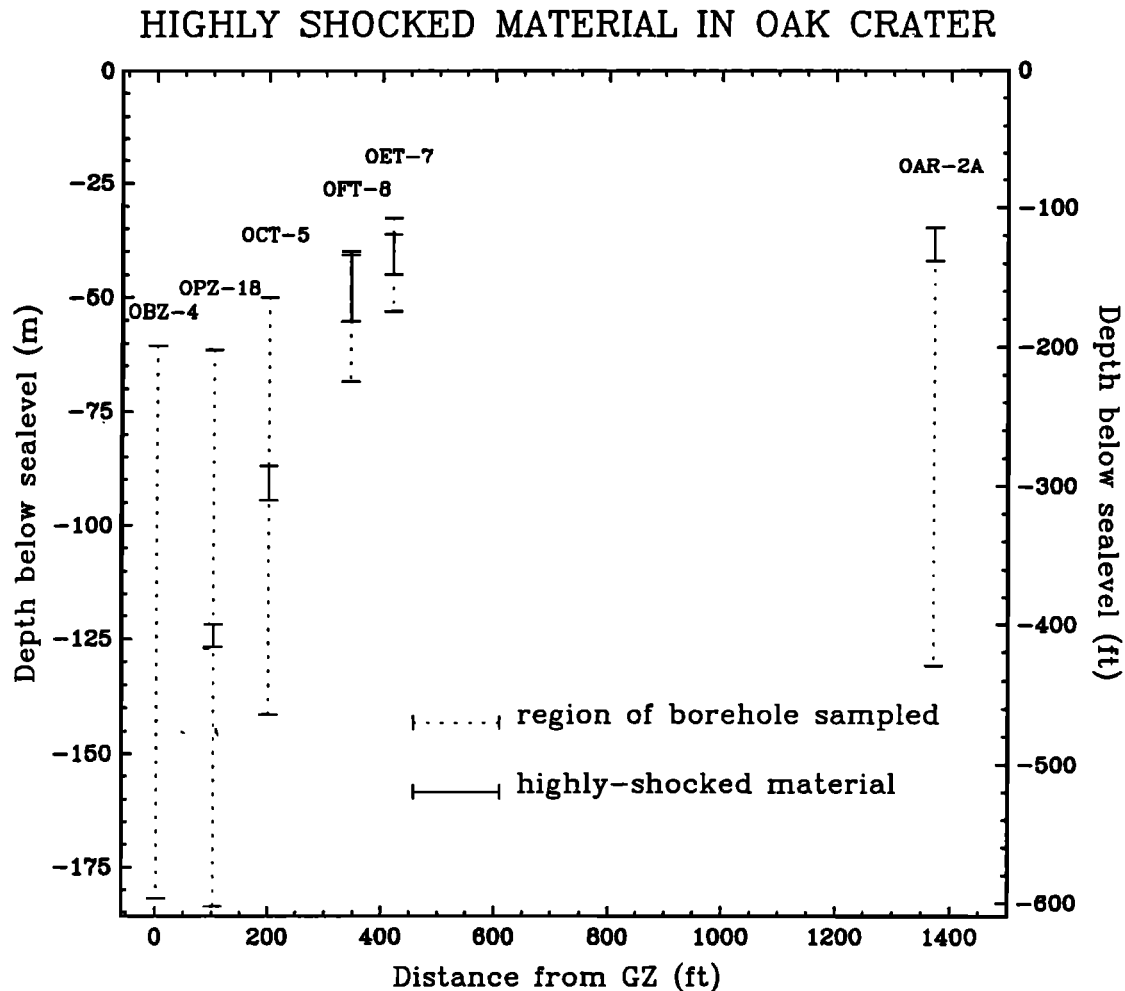


Figure 6a. Results of the borehole sample analysis showing shock pressure versus depth for each borehole.



**Figure 6b.** Results of the core sample analysis showing the depth and thickness of each zone of highly shocked material as a function of distance from ground zero. The dotted line indicates the distance from the current sea floor to the top of the zone, whereas the solid line indicates the region containing highly shocked material.

material most likely originated near the preshot seafloor surface. Three other sand samples above this layer, 117.9, 112.3, and 108.9 m bsl, were shocked to low levels (Figure 6a). The heavily shocked samples were located primarily in the geologic zone Beta 2, the transition sands, while the lightly shocked material came from the zone Beta 1 [Wardlaw and Henry, 1986]. The remaining 24 of the 31 samples appear to be unshocked. Remarkably, not one of the samples from OBZ-4 showed signs of significant shock damage. There were three samples from the Beta 1 zone that did register marginally detectable degrees of shock damage. Sufficient samples were analyzed from the transition sands and vicinity to characterize the core; therefore, it appears that OBZ-4 did not share the same history as OPZ-18.

Thick zones of highly shocked material were found in each of the three northeastern radial transition boreholes OCT-5, OET-7, and OFT-8. The transition sands have not been identified in any of these boreholes; however, the spectra of the shocked material are similar to those from the shocked material in OPZ-18.

Spectra were taken of 25 samples from borehole OCT-5. The results of six samples define a heavily shocked zone at least 7.5 m thick, extending from 87.0 to 94.4 m bsl. This region occurs within the Beta 1b (early-stage, collapse rubble) zone, and these samples are also primarily uncemented sands. Aside from the highly shocked material in this zone, there are four widely dispersed samples that appear to be moderately shocked. However, one sample in particular, taken at 141.4 m bsl, is an example of the aforementioned situation, where poor signal quality biases a pressure determination. Simple visual analysis of its spectra suggests that it is actually unshocked. The elevated pressure calculated for this depth is an artifact of the noisiness of the ESR spectra of the sample.

Borehole OFT-8 is located just within the transient crater [Henry *et al.*, 1986]. In this case, the region of heavily shocked material begins near the top of the Beta 1b zone and extends downward for approximately 8 m. Included within this zone were seven heavily shocked samples located between 46.8 and 55.1 m bsl.

**Table 3.** Results for OAK Ejecta Samples

Sample	Range, m	P, GPa	Source Depth, m bsl
78	321± 4	0.0± 0.9	—
125a	388± 4	0.0± 1.8	61–150
126	369± 4	13.6± 4.0	32–43
127	334± 4	3.0± 2.0	—
128	433± 4	10.5± 8.0	32–43
144	397± 4	12.5± 2.3	32–43
147	357± 9	2.7± 0.9	—
155	396± 4	0.0± 1.4	—
156	338± 4	0.0± 1.1	—
156b	338± 4	0.0± 0.9	—
158	276± 4	0.0± 1.0	61–150
166B	338± 4	0.0± 1.0	150–210
167B	389± 4	0.0± 0.6	—
168A	352± 4	0.0± 0.8	—
168C	342± 4	0.0± 1.4	91–150
201	704± 5	0.0± 1.4	130
1–1	414±48	13.1± 1.6	32–43
1–2	414±48	15.0± 3.4	>43
1–3	414±48	14.6± 1.6	32–43

Source depths are converted to meters below sea level from *Ludwig et al.* [1986] and B. Ristvet (personal communication, 1981).

Bordering this region above and below are zones containing moderately shocked material. The next farthest borehole from ground zero was OET-7. Based on seismic reflection, paleontology, and lithostratigraphic analysis data, this borehole is thought to be located outside of the transient crater [*Henry et al.*, 1986]. The majority of the samples were from the Gamma zone; however, all but the uppermost sample were heavily to moderately shocked. Of the highly shocked samples, six out of seven were uncemented sediment samples. The highly shocked zone extended from 36.2 to 45.0 m bsl.

Borehole OAR-2A was initially sampled only as a reference core; however, six surface samples appear to have been heavily shocked. All of the shocked samples were located within the top 12 m of the core, 33.9 to 45.4 m bsl, with the most heavily shocked material within the first 7 m. The proximity of this borehole to the reef suggests that highly shocked, fine-grain ejecta may have been deposited from the slope at later times. The combined results from the OAK borehole sample analysis are presented in Figure 6a. The solid horizontal line in each panel indicates the present seafloor depth. The depth and thickness of each zone containing highly shocked material ( $P = 10$  GPa) as a function of the distance of the borehole from ground zero are shown in a simplified manner in Figure 6b.

### Results of the Ejecta Sample Analysis

The OAK ejecta studied consist of 14 samples collected by a submersible from various sites throughout the crater and three samples collected by scuba divers from roughly a single site (Table 3). The former samples are a subset of a series of ejecta samples analyzed

by *Halley et al.* [1986]. Figure 7 shows the locations from where each ejecta sample was recovered. The range values were measured from the Figure 7 map. Unfortunately, the ejecta samples in this study were all taken from roughly the same distance from ground zero. There was only one sample, OAK 201, which was recovered at a significantly different range. The results of the ejecta analysis are plotted in Figure 8a. The majority of the ejecta samples were relatively unshocked; however, all of the highly shocked ejecta were found at the base of the reef slope. It is highly unlikely that the reef structure would have blocked the ejection of the highly shocked material during the formation of the transient crater, as the initial ejection angles near surface ground zero are calculationally and experimentally verified to be high angle ( $>40^\circ$ ). Underwater photos reported by *Tremba et al.* [1982] and *Shinn et al.* [1986] strongly support the hypothesis that much of the ejecta was transported down the reef slope during transient crater wall collapse, or later, during the formation of the present crater.

In addition to the range measurements, the estimated preexplosion initial depth of a limited number of the ejecta samples was available from strontium isotopic analyses [*Halley et al.*, 1986] and paleontology (B. Ristvet, personal communication, 1981). The preexplosion depth below seafloor is plotted against shock pressure for these samples in Figure 8b. Although the preexplosion depth estimates are crude, there is a strong correlation between shock pressure and depth for this limited data set. This is consistent, however, with the previous assertion that the surface material was the most severely shocked.

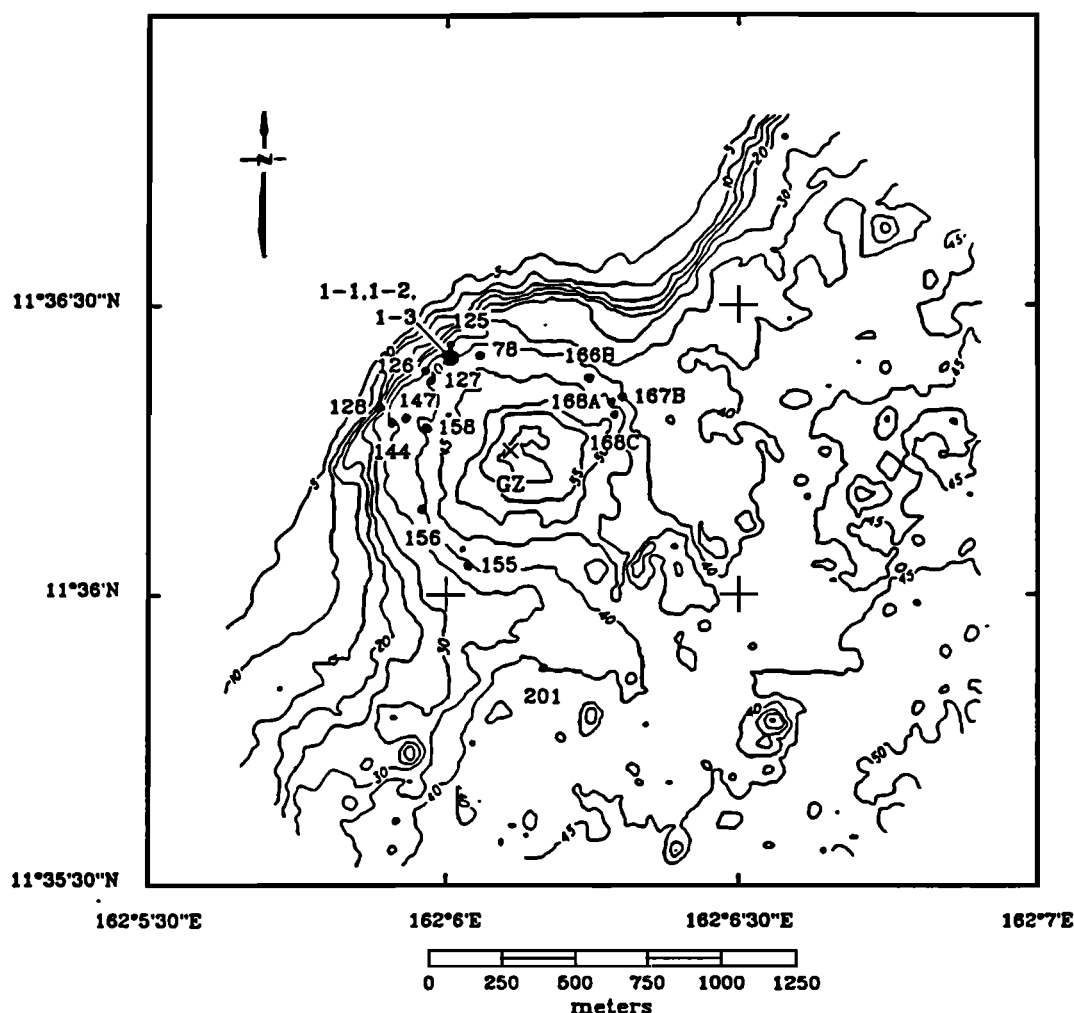


Figure 7. Map of OAK crater showing ejecta recovery sites. Origin of map same as Figure 5.

### Interpretation of OAK Crater

With a few exceptions, the bulk of the samples analyzed can be split into two categories, unshocked and very heavily shocked. There were relatively few samples that are assigned to intermediate pressures. The lack of samples at intermediate shock pressures suggests that the majority of the shocked material shares a common origin, which appears to be initially near-surface material which outlines the original transient crater. Below this heavily shocked material, the rock core samples, although deformed, are in proper stratigraphic sequence. Above the heavily shocked zone the sediment and rock in the crater appear to be fallout ejecta material introduced into the crater after the explosion via shock-induced liquefaction [Tremba *et al.*, 1990], and subsequent wave and storm sediment transport.

Assuming the highly shocked material shown in Figure 6 represents the outline of the present shape of the transient cavity and this cavity is circularly symmetric, a volume of  $3.6 \pm 0.2 \times 10^6 \text{ m}^3$  is calculated from the data of Figure 6. Because of postexplosion upward deformation, this volume should be considered a lower bound. Notably, Tremba *et al.* [1990] on the basis of

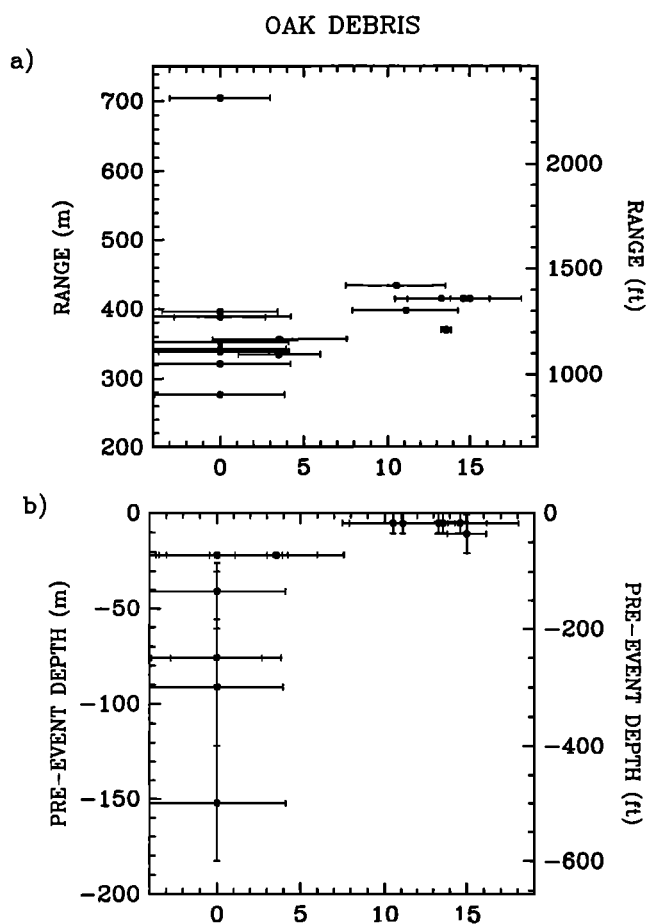
this and other geologic data, inferred a transient crater with nearly twice this volume of  $6.8 \times 10^6 \text{ m}^3$ . Schmidt *et al.* [1986] have developed a scaling relation for both the volume of the final crater and transient cavity versus energy for surface explosions in water-saturated rock. Here, in cgs units,

$$(V/E) = 5.2 \times 10^{-7} E^{-0.2} \quad (4)$$

Equation (4) indicates that the volume excavated ( $V$ ) divided by explosive energy,  $E$ , decreases with increasing energy. Small-scale model explosive experiments conducted under high gravities generated in the Boeing centrifuge facility indicate that the volume of the transient cavity is 1.4 times that of the final crater. Equation (4) yields a final crater volume of  $3.7 \times 10^6 \text{ m}^3$  and hence, a transient crater volume (a factor of 1.4 greater) of  $5.2 \times 10^6 \text{ m}^3$ , in good agreement with the Tremba *et al.* [1990] value.

### Meteor Crater

Meteor Crater is 180 m deep and 1.2 km in diameter. It is located in north-central Arizona on the Colorado



**Figure 8.** Results of the OAK Crater ejecta sample analysis showing (a) shock pressure as a function of range from ground zero, and (b) estimated preexplosion depth below seafloor versus shock pressure for a subset of the ejecta samples.

Plateau. The crater was formed in flat-lying bedrock consisting of the Moenkopi, Kaibab, Coconino, and Toroweap formations [Shoemaker and Kieffer, 1974]. Thermoluminescence studies of the Coconino sandstone and the Kaibab dolomite give an age of 49,000 years for the crater [Sutton, 1985].

Unlike the previous discussion of OAK Crater, the analysis of the Meteor Crater samples is less complete. The Kaibab Formation data set is divided into five sections: pressure calibration experiments (discussed above), Diablo Canyon unshocked Kaibab standards, samples from the walls of Meteor Crater, Meteor Crater ejecta samples, and a miscellaneous category including

breccia samples and highly shocked Kaibab Formation ejecta subsequently redeposited inside the crater. A complete catalog of these spectra is given by Polanskey [1989].

Although the effects of shock pressure on Kaibab Formation samples have been demonstrated in Figure 2b, shock pressure may not be the only variable. As mentioned previously, the  $\text{Mn}^{2+}$  may substitute in both the Mg and Ca sites of the dolomite structure. The distribution of  $\text{Mn}^{2+}$  between these sites will affect the shape of the resulting spectrum. Therefore, differences in the  $\text{Mn}^{2+}$  distribution between the various members of the Kaibab Formation could confuse the effects of shock pressure. The spectra from a series of Kaibab samples taken from Diablo Canyon (16 km east of Meteor Crater) are presented by Polanskey [1989]. Although the calibration equations (3a) and (3b) were defined using the high-resolution spectra constructed from an average of the Diablo Canyon samples, the calibration was still applied to aliquots of the unshocked samples as a self-consistency test. The results are listed in Table 4 and show that in most cases the expected zero pressure was obtained. The pressure of 0.1 GPa calculated for sample 2 illustrates the limitations in accuracy of this analysis. A sample of caliche was also processed in order to determine whether or not weathering products would influence the analysis. The caliche spectrum was clearly distinguishable from the Kaibab spectra.

## Results of the Meteor Crater Sample Analysis

The results for the crater wall analysis are listed in Table 5. There is a slight indication of shock damage in the 0.3 to 0.6 GPa range in the samples from the  $\beta$  member of the Kaibab formation.

The peak pressure range experienced by some Kaibab limestone samples from the crater wall may be used to place constraints on the energy of the impactor using the relations derived by Moss [1988] and Lamb *et al.* [1991] for free-field shock pressure versus radius. Since we have in situ sampled rocks which lie not beneath the crater, which is less affected by the free surface, but a zone close to the free surface which may have experienced less than the free field shock pressure on account of stress wave reflections at the free surface [Melosh, 1984], lower bounds on the impactor energy can only be obtained.

The Moss-Lamb model assumes that for an explosion in the free field the total energy,  $W$ , is uniformly de-

**Table 4.** Results for Diablo Canyon Samples

Sample Number	Pressure, GPa	Description
1	$0.0 \pm 0.03$	mid- $\alpha$ member, near surface
2	$0.1 \pm 0.01$	lower- $\alpha$ member
3	$0.0 \pm 0.05$	upper- $\beta$ member, near contact to $\alpha$ member
4	$0.0 \pm 0.01$	mid- $\beta$ member, 30 ft (10 m) into member
5	$0.0 \pm 0.03$	60-70 ft (18-21 m) into $\beta$ member

**Table 5.** Results for Meteor Crater Wall Samples

Sample Number,	Pressure, GPA	Description
1	0.0 ± 0.02	uppermost $\alpha$ member, contact Kaibab/Moenkopi
2	0.0 ± 0.02	upper $\alpha$ member, below yellow vuggy dolomite
3	0.1 ± 0.02	lower $\alpha$ member, below breccia
4	0.3 ± 0.04	uppermost $\beta$ member, above promontory
5	0.5 ± 0.06	upper $\beta$ member, above promontory
6	0.6 ± 0.08	lower $\beta$ member, near promontory
7	0.4 ± 0.05	lower $\beta$ member, near Astronaut trail
8	0.1 ± 0.02	$\gamma$ member, base of promontory

posited within a spherical shock of radius,  $R$ , and the internal energy and kinetic energy per unit mass are equal; then

$$fW = u^2 (\rho_o \frac{4}{3} \pi R^3) \quad (5)$$

Here  $f$  is an efficiency factor which takes into account that because of irreversible energy deposition only a fraction of the energy is available to drive a shock. Also,  $u$  is the outward particle velocity of the shock wave, and  $\rho_o$  is the initial medium density. The term within the parentheses is the mass encompassed by the shock and  $f$  would be unity if the work done on the rock within the sphere was all reversible. By fitting the observed decays of the stress waves from a large number of nuclear explosions, both in wet and dry rocks, Moss found that  $f \approx 0.53$ . For many dry and wet media the shock velocity versus particle velocity is given as

$$U_s = C_o + su \quad (6)$$

Here  $C_o$  and  $s$  are the parameters for Kaibab limestone of Table 1. The shock pressure is given by the Rankine Hugoniot equation

$$P = \rho_o U_s u \quad (7)$$

Upon substituting  $U_s$  from (6) in (7) and then eliminating  $u$ , between (5) and (7), we solve the resulting expression for  $W$  using the 0.53 value for  $f$ . Assuming a peak value of pressure (Table 4) from samples taken at a 0.533-km radius from the crater center of  $0.6 \pm 0.08$  GPA yields an impactor energy,  $W$ , of  $8.9 \pm 2.3$  Mt, whereas a lower value of  $0.3 \pm 0.04$  GPA yields an impact energy of  $2.41 \pm 0.65$  Mt. We use the greater of these bounds to test the *Schmidt and Housen* [1987] gravity cratering scaling law for wet granular media. This is given by

$$\Pi_r = 0.8 \Pi_2^{-0.22} \quad (8)$$

where

$$\Pi_r = R (\rho_o/m)^{1/3} \quad (9)$$

where  $m$  is the mass of the impactor,  $\rho_o$  is the initial density of the medium, and  $R$  is the crater radius. Also

$$\Pi_r = 3.22 g a/U^2 \quad (10)$$

where  $g$  is gravity,  $a$  and  $U$  are the projectile radius and impact velocity. Numerical calculational fits of the theoretical crater to the observed crater have assumed values of  $51 \times 10^9$  to  $167 \times 10^9$  g [e.g., *Roddy et al.*, 1980; *Bryan et al.*, 1978] for an assumed iron-nickel, density  $7.86$  g/cm<sup>3</sup> bolide. Using a crater radius of  $0.533$  km, and an average water-saturated target density of  $2.35$  g/cm<sup>3</sup> in (8)-(10) yields an impactor energy of  $4.06$  and  $3.30$  Mt, over the range of bolide masses assumed. These energy estimates are in good agreement but somewhat lower than our maximum energy of  $8.9 \pm 2.3$  Mt inferred from the peak shock pressure seen by rock in the crater wall.

Another approach to using the present available theory of *Schmidt and Housen* is to employ the calculation of crater volume from *Roddy* [1978].

Analogous to (9) is:

$$\Pi_2 = \rho V/m \quad (11)$$

where  $\rho$  is the density of the target ( $2.35$  g/cm<sup>3</sup>),  $V$  is *Roddy's* [1978] value of  $75 \times 10^{12}$  cm<sup>3</sup> for the "apparent crater volume," and  $m$  is the assumed bolide mass, here again taken to range from  $51$  to  $167 \times 10^9$  gm. Using the definition of  $\Pi_v$  as

$$\Pi_v = 3.22ga/U^2 \quad (12)$$

and the *Schmidt and Housen* [1987] equation

$$\Pi_v = 0.2\Pi_2^{-0.65} \quad (13)$$

we solve (11), (12), and (13) for  $U$  (impact velocity) for the smaller and larger assumed iron-nickel bolide mass to obtain  $34.8$  and  $17.0$  km/s. These impact velocities then yield impact energies of  $7.4$  and  $5.8$  Mt, respectively. These are somewhat larger energies than originally estimated by *Shoemaker* [1963] (of  $1.2$  to  $1.8$  Mt) on the basis of cube-root scaling from the nuclear craters.

Finally, we note there was also some evidence of light shock damage in the ejecta samples (Table 6). Figure 9 shows a plot of shock pressure versus distance from the crater rim for the Meteor Crater ejecta. As in the case of OAK Crater ejecta, the majority of the samples appear to be unshocked. However, there was some bias in the ejecta sampling technique. All the ejecta samples

**Table 6.** Results for Meteor Crater Ejecta Samples

Sample	Pressure, GPa	Range, m
1	0.0± 0.04	1414
2	0.3± 0.04	1240
3	0.0± 0.07	1240
4	0.5± 0.08	1212
5	0.1± 0.02	1183
6	0.0± 0.07	1178
7	0.0± 0.06	1082
8A	0.0± 0.06	928
8B	0.1± 0.04	928
9	0.0± 0.01	350
10	0.4± 0.08	284
11	0.6± 0.05	140

were taken from large boulders within the ejecta field. The results could change if more samples were taken over a complete distribution in ejecta size. No sign of the  $Mn^{2+}$  spectrum was found for ejecta samples which had been partially melted as a result of the impact. In this material the carbonate phases were probably not present.

## Summary

The comparison of EPR spectra from samples shocked in the laboratory with those subjected to impact and explosion-induced shocks has been shown to be a useful technique in the analysis of craters in carbonate rocks. In OAK Crater we were able to detect for the first time, the highly shocked (10-13 GPa) lining

of the transient crater beneath the apparent crater. It should be emphasized that this material has undergone substantial uplift and other motion since it lined the transient crater.

The volume of this lining,  $3.2 \pm 0.2 \times 10^6 \text{ m}^3$  at a depth of 125 m below sea level beneath a 60 m cover of fallback ejecta and other material at the center of the crater provides a test of the scaling for transient craters proposed by *Schmidt et al.* [1986]. They propose on the basis of small-scale laboratory experimentation, a relation for the transient cavity produced by a surface explosion of

$$V = 7.3 \times 10^{-7} E^{0.8} \quad (\text{cgs units}) \quad (4')$$

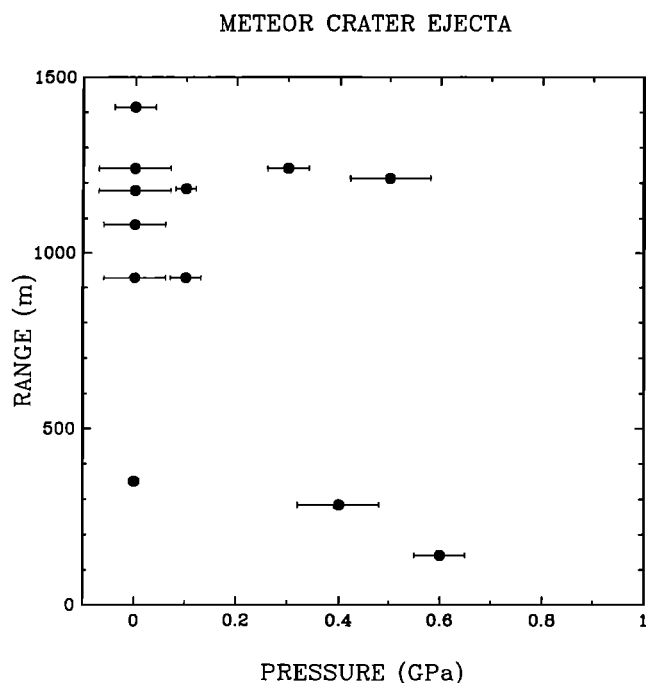
For a known yield, (4') predicts a volume of  $5.3 \times 10^6 \text{ m}^3$ . Since our experimental estimate of transient crater volume,  $3.2 \pm 0.2 \times 10^6 \text{ m}^3$  is a lower bound, we consider the agreement good. The occurrence of heavily shocked (10-15 GPa) ejecta, which had an initial shallow depth in the carbonate reef material of only 32 to 43 m, is in general accord with crater models which predict that initial surface materials should be more heavily shocked.

The demonstration that the carbonate rocks of the 49,000-year-old Barringer, Arizona (Meteor), Crater retain the shock-induced internal deformation signature suggests that ESR methods can be utilized in the study of geologically young craters in carbonate terranes.

Kaibab limestone ejecta cobbles and carbonate in situ in the walls of the crater demonstrate the effects of shock deformation to stress levels of 0.3 to 0.6 GPa. Using this peak shock pressure and the *Moss* [1988] and *Lamb et al.* [1991] relations for shock pressure versus radius for explosion-induced shock wave decay in the free field of  $2.41 \pm 0.65$  and  $8.9 \pm 2.3 \text{ Mt}$  is inferred for the bolide energy required to produce Meteor Crater. This value is comparable to those obtained from the *Schmidt and Housen* [1987] scaling law which, depending on assumptions regarding the bolide mass, yield impact energies of 3.3 to 4.1 Mt for crater radius scaling. If instead we use the inferred apparent true crater volume of *Roddy* [1978], their scaling relations imply an energy of 5.8 to 7.4 Mt.

Thus, we conclude that in the case of Meteor Crater, the peak shock pressure seen by in situ carbonate rocks on the rim (0.3 to 0.6 GPa) give impact energies (2.4-8.9 Mt), which are in good agreement derived from modern theories (3.3. to 7.4 Mt) of the energetics of impact cratering.

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**Figure 9.** Results of the Meteor Crater ejecta sample analysis showing shock pressure as a function of range from the crater rim.



girda, Byron Ristvet, David Live, and Fun-dow Tsai were most helpful. Division of Geological and Planetary Sciences, California Institute of Technology, Pasadena, contribution 4704.

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